Table 1 Masses of Double Neutron Star Systems and Non-recycled Pulsars

System	M_T ($M_\odot$)	M_{PSR} ($M_\odot$)	M_c ($M_\odot$)	Mass const.	Ref.
Systems with well-measured component masses					
J0453+1559	2.734(4)	1.559(5)	1.174(4)	$\dot{\omega}, h_3$	1,2
J0737−3039	2.58708(16)	1.3381(7)	1.2489(7) y	$\dot{\omega}, q$	3
B1534+12	2.678463(8)	1.3330(4)	1.3455(4)	$\dot{\omega}, \gamma$	4
J1756−2251	2.56999(6)	1.341(7)	1.230(7)	$\dot{\omega}, \gamma$	5
J1906+0746	2.6134(3)	1.291(11) y	1.322(11) ?	$\dot{\omega}, \gamma$	6
B1913+16	2.828378(7)	1.4398(2)	1.3886(2)	$\dot{\omega}, \gamma$	7
B2127+11C g	2.71279(13)	1.358(10)	1.354(10)	$\dot{\omega}, \gamma$	8
Systems with total binary mass measurement only					
J1518+4904	2.7183(7)	<1.768	>0.950	$\dot{\omega}$	9
J1811−1736	2.57(10)	<1.64	>0.93	$\dot{\omega}$	10
J1829+2456	2.59(2)	< 1.34	>1.26	$\dot{\omega}$	11
J1930−1852	2.59(4)	< 1.32	>1.30	$\dot{\omega}$	12
Non-recycled pulsars with massive WD companions					
J1141−6545	2.2892(3)	1.27(1) y	1.01(1)	$\dot{\omega}, \gamma$	14,15
B2303+46	2.64(5)	1.24-1.44 y	1.4-1.2	$\dot{\omega}, M_{WD}$	15,16

Notes: The systems indicated with a "g" are located in globular clusters. A question mark indicates that the NS nature of the companion is not firmly established. The mass measurements for neutron stars detected as normal (non-recycled) radio pulsars are indicated with the letter "y". References are to the latest mass measurements: 1. Deneva et al. (2013) 2. Martinez et al. (2015) 3. Kramer et al. (2006) 4. Fonseca, Stairs & Thorsett (2014) 5. Ferdman et al. (2014) 6. van Leeuwen et al. (2015) 7. Weisberg, Nice & Taylor (2010) 8. Jacoby et al. (2006) 9. Janssen et al. (2008) 10. Corongiu et al. (2007) 11. Champion et al. (2005) 12. Swiggum et al. (2015) 13. Bhat, Bailes & Verbiest (2008) 14. Antoniadis et al. (2011) 15. Thorsett & Chakrabarty (1999) 16. van Kerkwijk & Kulkarni (1999)

Table 2 Masses of Millisecond Pulsars

System	M_T ($M_\odot$)	M_{PSR} ($M_\odot$)	M_c ($M_\odot$)	Mass const.	Ref.
MSPs with WD companions and low-eccentricity orbits					
J0348+0432		2.01(4)	0.172(3)	q, M_{WD}	Antoniadis et al. (2013)
J0437−4715		1.44(7)	0.224(7)	r, s	Reardon et al. (2016)
J0621+1002	2.32(8)	$1.53^{+0.10}_{-0.20}$	$0.76^{+0.28}_{-0.07}$	$\dot{\omega}, s$	Kasian (2012)
J0751+1807		1.72(7)	0.13(2)	$s, \dot{P}_b$	Desvignes (2015)
J1012+5307		1.83(11)	0.16(2)	q, M_{WD}	Antoniadis (2015)
J1614−2230		1.928(7)	0.500(6)	r, s	Arzoumanian et al. (2015)
J1713+0747		1.31(11)	0.286(12)	r, s	Zhu et al. (2015)
J1738+0333		$1.47^{+0.07}_{-0.06}$	$0.181^{+0.007}_{-0.005}$	q, M_{WD}	Antoniadis et al. (2012)
J1802−2124		1.24(11)	0.78(4)	r, s	Ferdman et al. (2010)
J1807−2500B	2.57190(73)	1.3655(21)	1.2064(20)(?)	$\dot{\omega}, h_3$	Lynch et al. (2012)
B1855+09		1.58^{+10}_{-13}	$0.267^{+0.010}_{-0.014}$	r, s	Splaver (2004)
J1909−3744		1.47(3)	0.2067(19)	r, s	Reardon et al. (2016)
J2222−0137		1.20(14)	1.05(6)	r, s	Kaplan et al. (2014a)
MSPs with eccentric orbits and triples					
J0337+1715		1.4378(13)	0.19751(15) 0.4101(3)	i, q	Ransom et al. (2014); Kaplan et al. (2014b)
J1903+0327	2.697(29)	1.667(21)	1.029(8)	$\dot{\omega}, h_3$	Freire et al. (2011)
J1946+3417	2.186(11)	1.912(10)	0.2734(11)	$\dot{\omega}, h_3$	Barr et al. (2015)
J2234+0611	1.668(6)	1.393(13)	0.276(9)	$\dot{\omega}, h_3$	Stovall & et al. (2015)
MSPs in globular clusters					
J0024−7204H	1.61(4)	< 1.52	> 0.164	$\dot{\omega}$	Freire et al. (2003)
J0514−4002A	2.453(14)	< 1.50	> 0.96	$\dot{\omega}$	Freire, Ransom & Gupta (2007)
B1516+02B	2.29(17)	< 2.52	> 0.13	$\dot{\omega}$	Freire et al. (2008a)
J1748−2021B	2.92(20)	< 3.24	> 0.11	$\dot{\omega}$	Freire et al. (2008b)
J1748−2446I	2.17(2)	< 1.96	> 0.24	$\dot{\omega}$	Ransom et al. (2005)
J1748−2446J	2.20(4)	< 1.96	> 0.38	$\dot{\omega}$	Ransom et al. (2005)
J1750−37A	1.97(15)	< 1.65	> 0.53	$\dot{\omega}$	Freire et al. (2008b)
B1802−07	1.62(7)	< 1.7	> 0.23	$\dot{\omega}$	Thorsett & Chakrabarty (1999)
J1824−2452C	1.616(7)	< 1.35	> 0.26	$\dot{\omega}$	Bégin (2006)
J1910-5958A		1.3(2)	0.180(18)	q, M_{WD}	Bassa et al. (2006); Cocozza et al. (2006)

Notes: J1807−2500B is located in the globular cluster NGC6544. A question mark indicates that the nature of the companion is uncertain. The total mass is indicated only when it is known more precisely than the masses of the components. References are to the latest mass measurements.

Table 3 Masses of Neutron Stars in High-Mass and Low-Mass X-ray Binaries

System	$M_{\rm NS}$ ($M_\odot$)	Error ($M_\odot$)	References
Neutron Stars in High-Mass X-ray Binaries			
LMC X−4	1.57	0.11	1
Cen X−3	1.57	0.16	1
4U 1538−522	1.02	0.17	1
SMC X−1	1.21	0.12	1
SAX J1802.7−2017	1.57	0.25	1
XTE J1855−026	1.41	0.24	1
Vela X−1	2.12	0.16	1
EXO 1722−363	1.91	0.45	1
OAO 1657−415	1.74	0.30	1
Her X−1	1.07	0.36	1
Neutron Stars in Low-Mass X-ray Binaries			
4U 1608−52	1.57	$^{+0.30}_{-0.29}$	2
4U 1724−207	1.81	$^{+0.25}_{-0.37}$	2
KS 1731−260	1.61	$^{+0.35}_{-0.37}$	2
EXO 1745−248	1.65	$^{+0.21}_{-0.31}$	2
SAX J1748.9−2021	1.81	$^{+0.25}_{-0.37}$	2
4U 1820−30	1.77	$^{+0.25}_{-0.28}$	2
Cyg X−2	1.90	$^{+0.22}_{-0.35}$	3

Notes. 1. See Falanga et al. (2015), Özel et al. (2012), Rawls et al. (2011), and references therein. We exclude 4U 1700−377, for which there is no evidence that it is a neutron star. 2. See Özel et al. (2015) for the latest constraints. 3. Orosz & Kuulkers (1999).